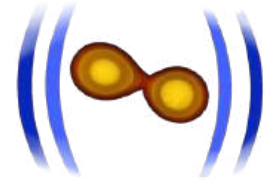




PennState
Eberly College of Science



CoRe collaboration

Neutron Rich Outflows from NS Mergers and AICs

David Radice – March 26, 2026

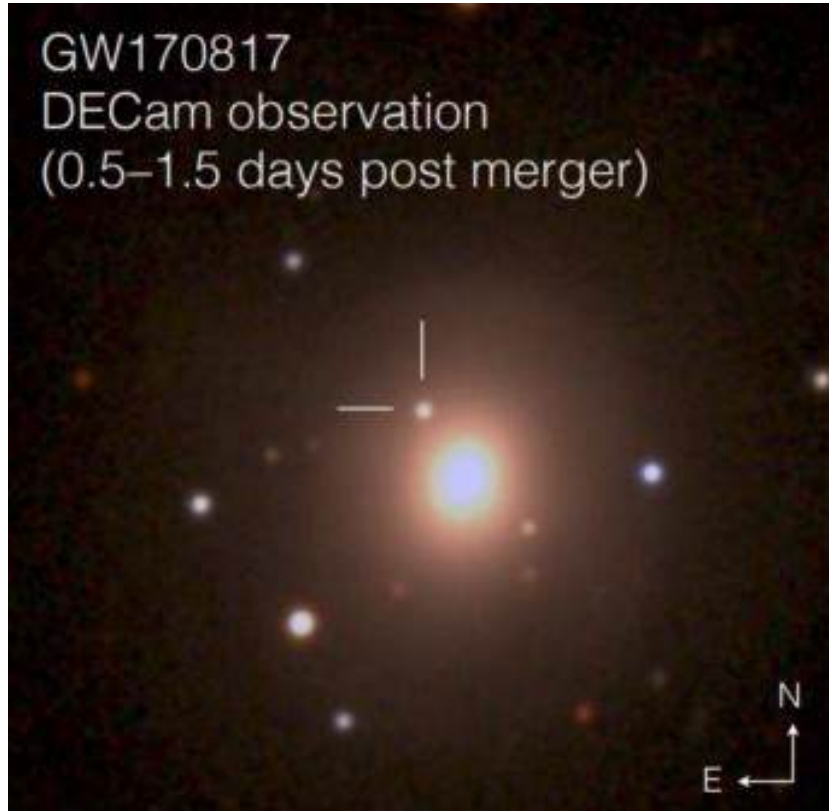


Network for Neutrinos,
Nuclear Astrophysics,
and Symmetries

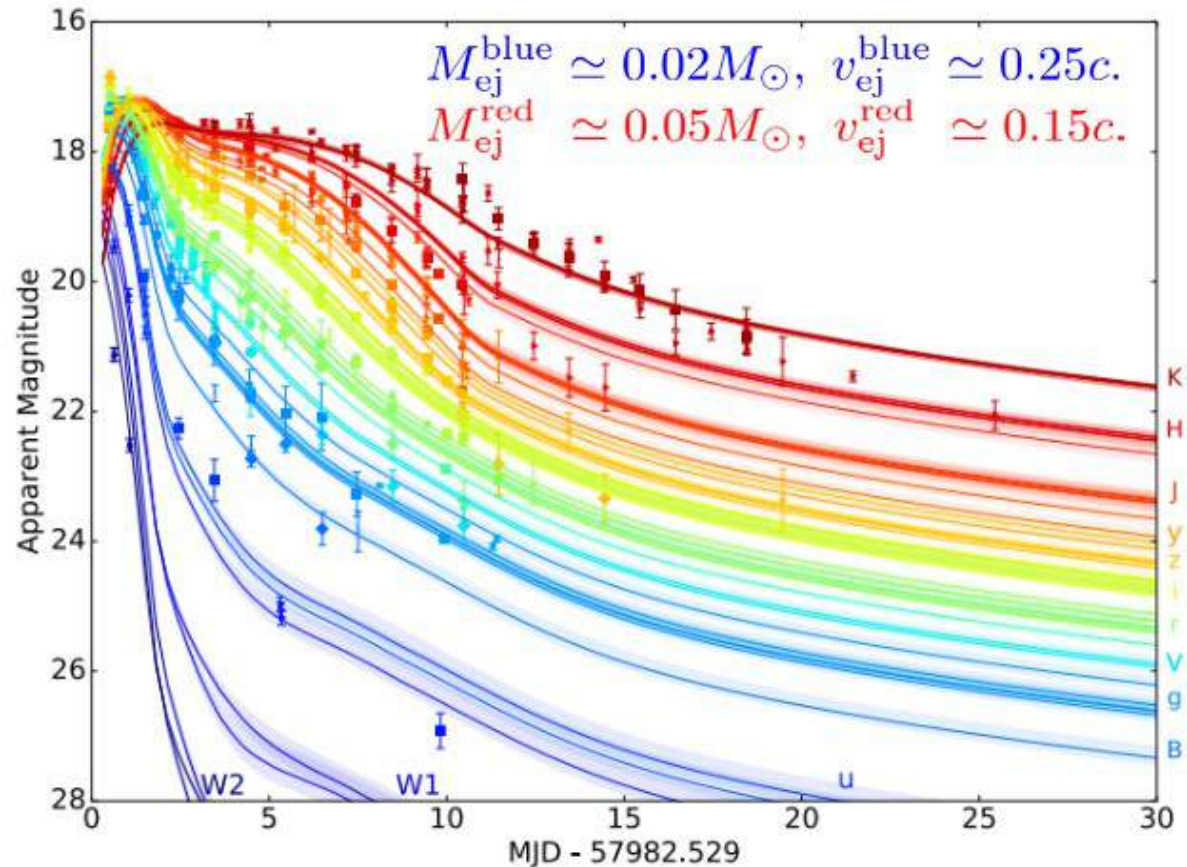
PHYSICS FRONTIER CENTER



R-process in neutron star mergers

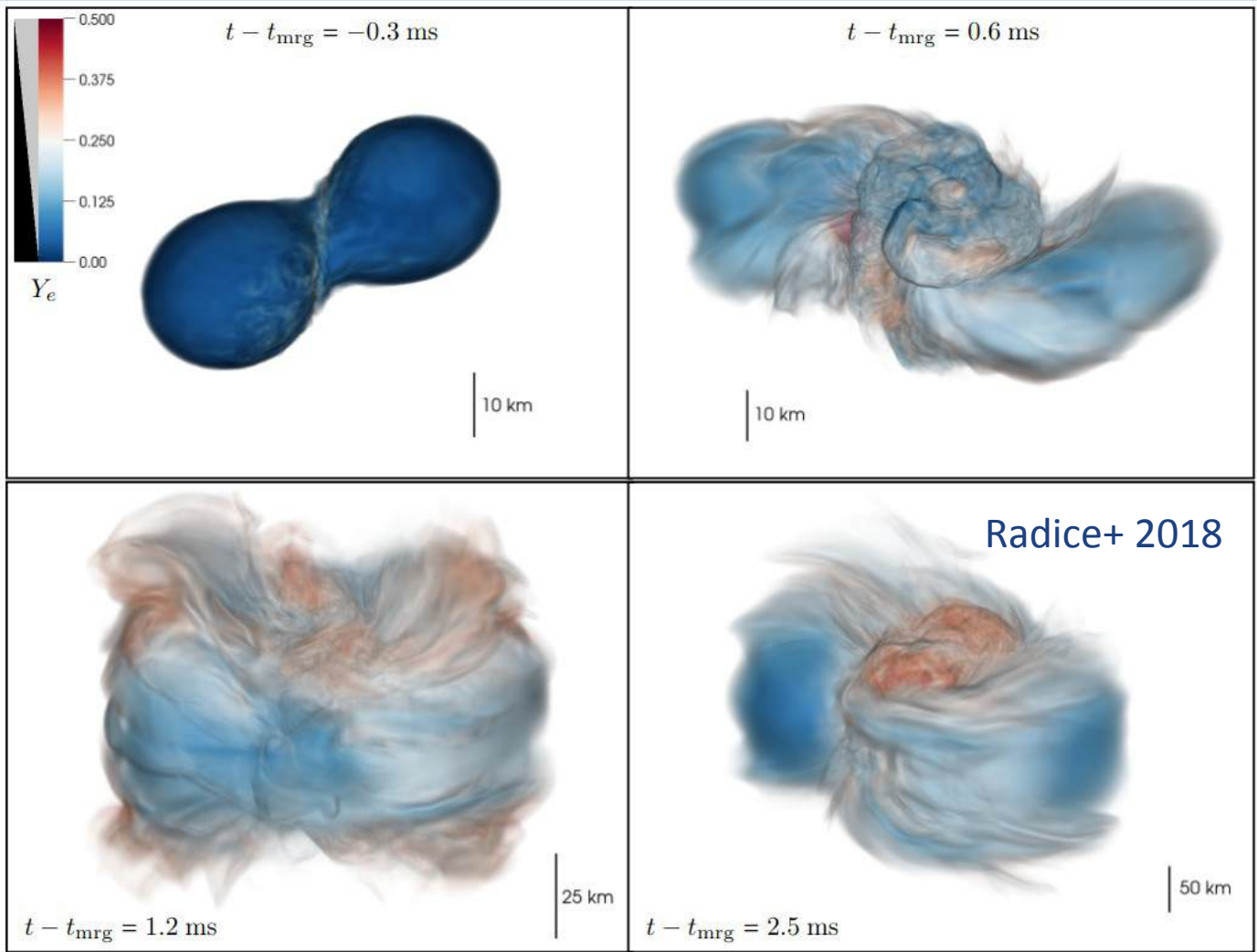


R-process in neutron star mergers

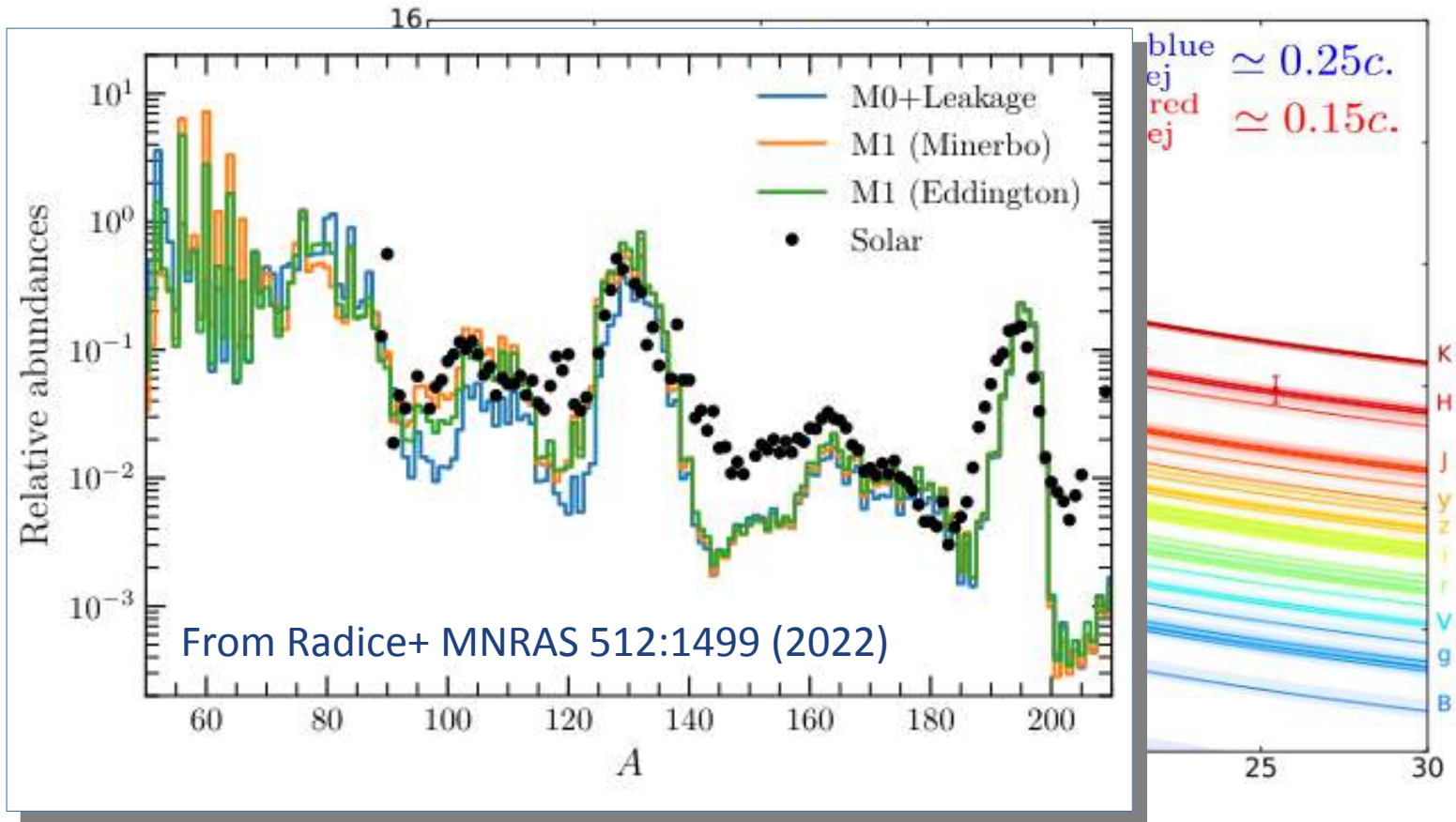


R-p

ers

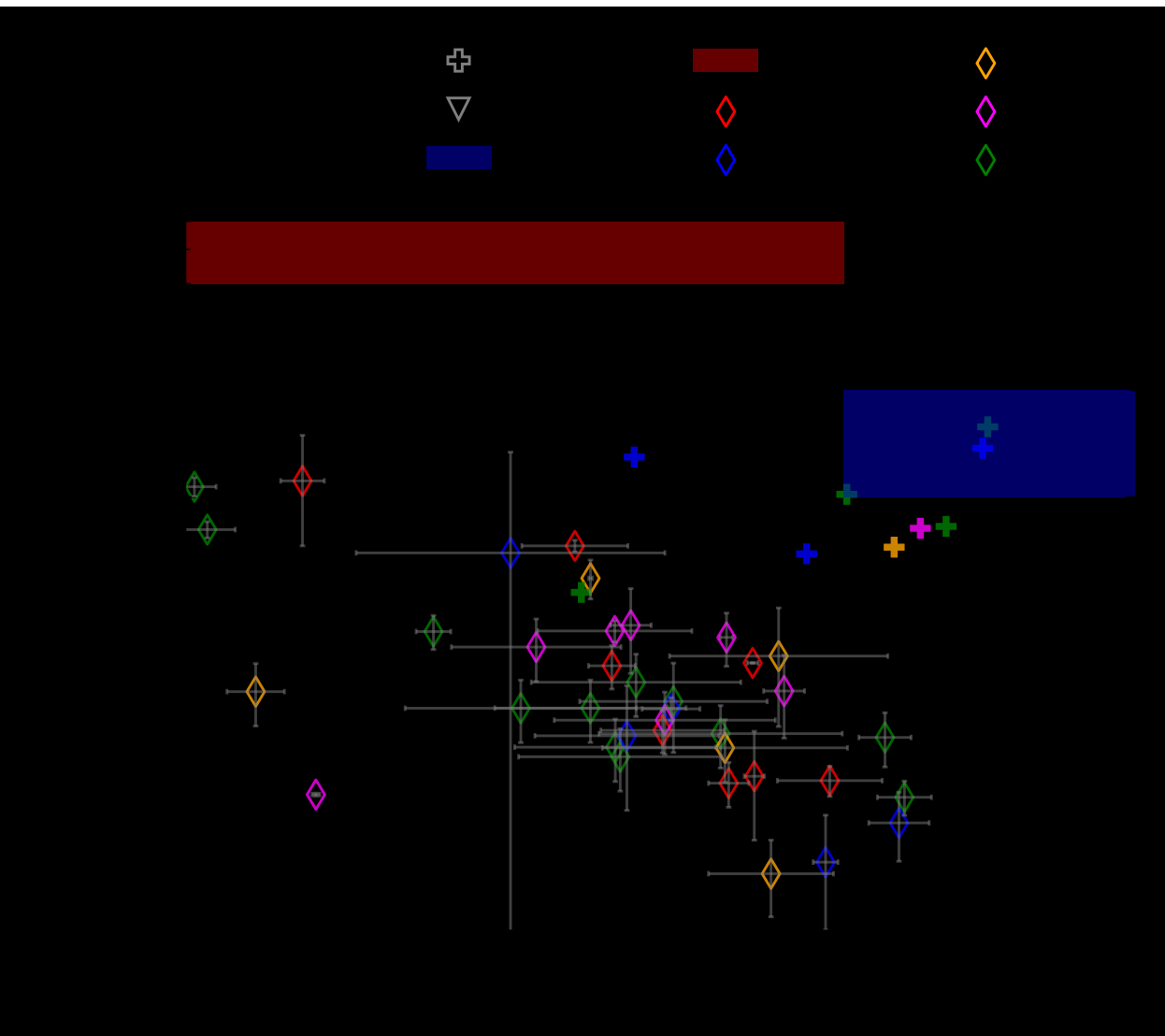


R-process in neutron star mergers



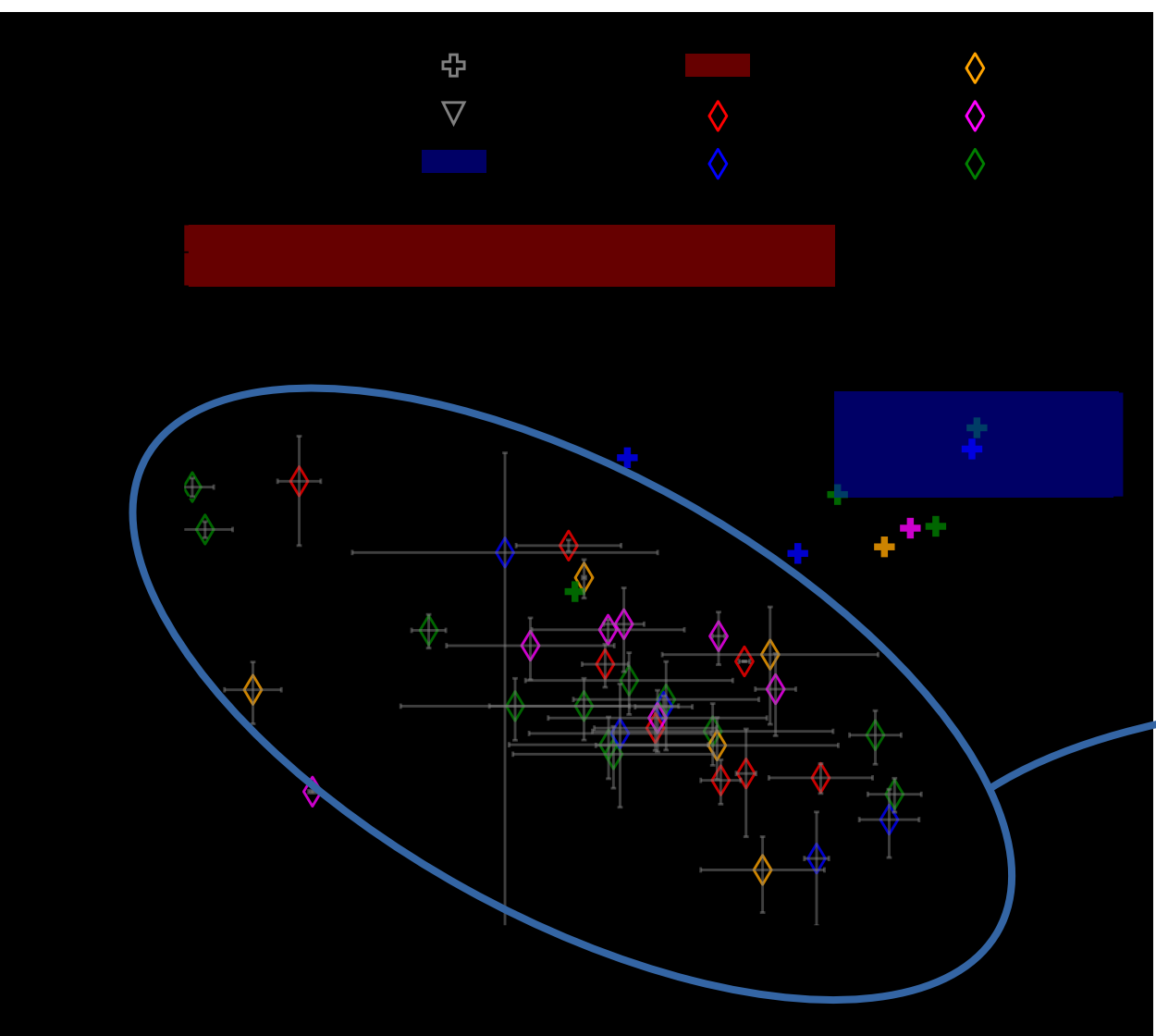
Problem solved?

- NS mergers make r-process elements, but in what **quantity** and **quality**?
Are they the only source?
- What about low-/high-metallicity trends, dwarf galaxies, ever decreasing LIGO rates (😞), etc.?
- Theory models are **incomplete** and affected by **systematic uncertainties**



See also Dessart+ 2009; Fernandez+ 2013; Perego+ 2014, Just+ 2014; ...

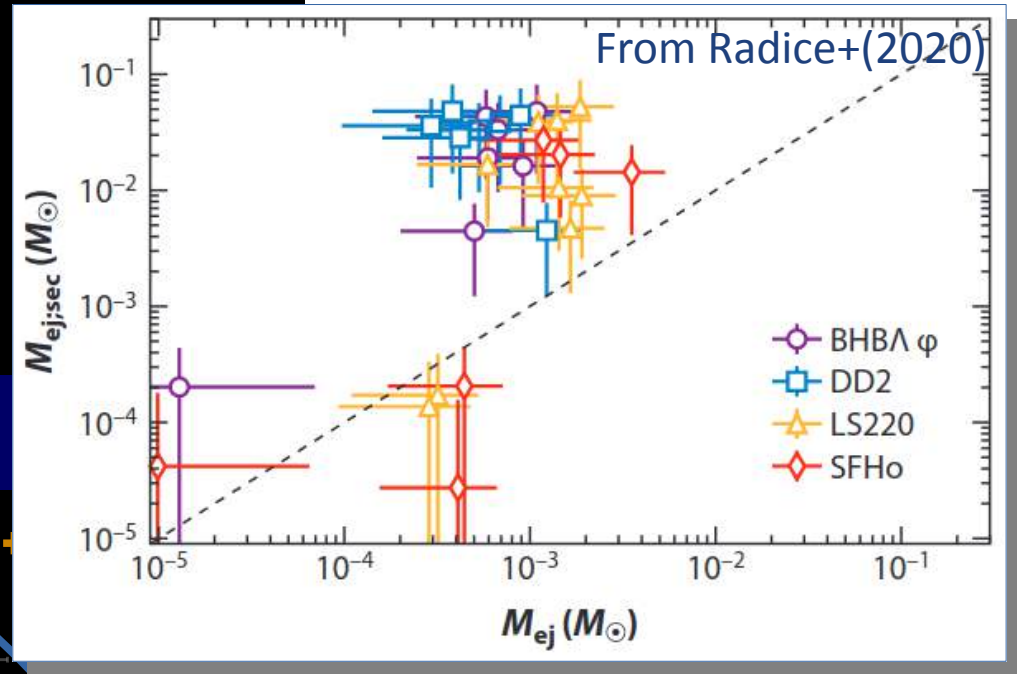
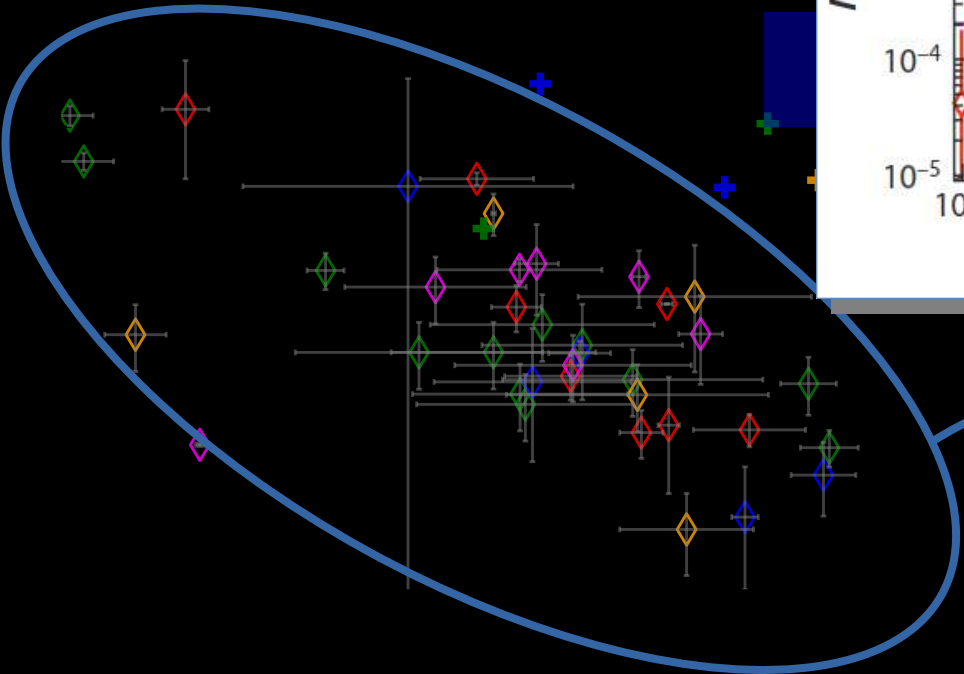
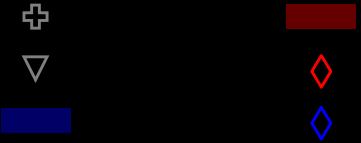
From Nedora+, ApJ 906:98 (2021)



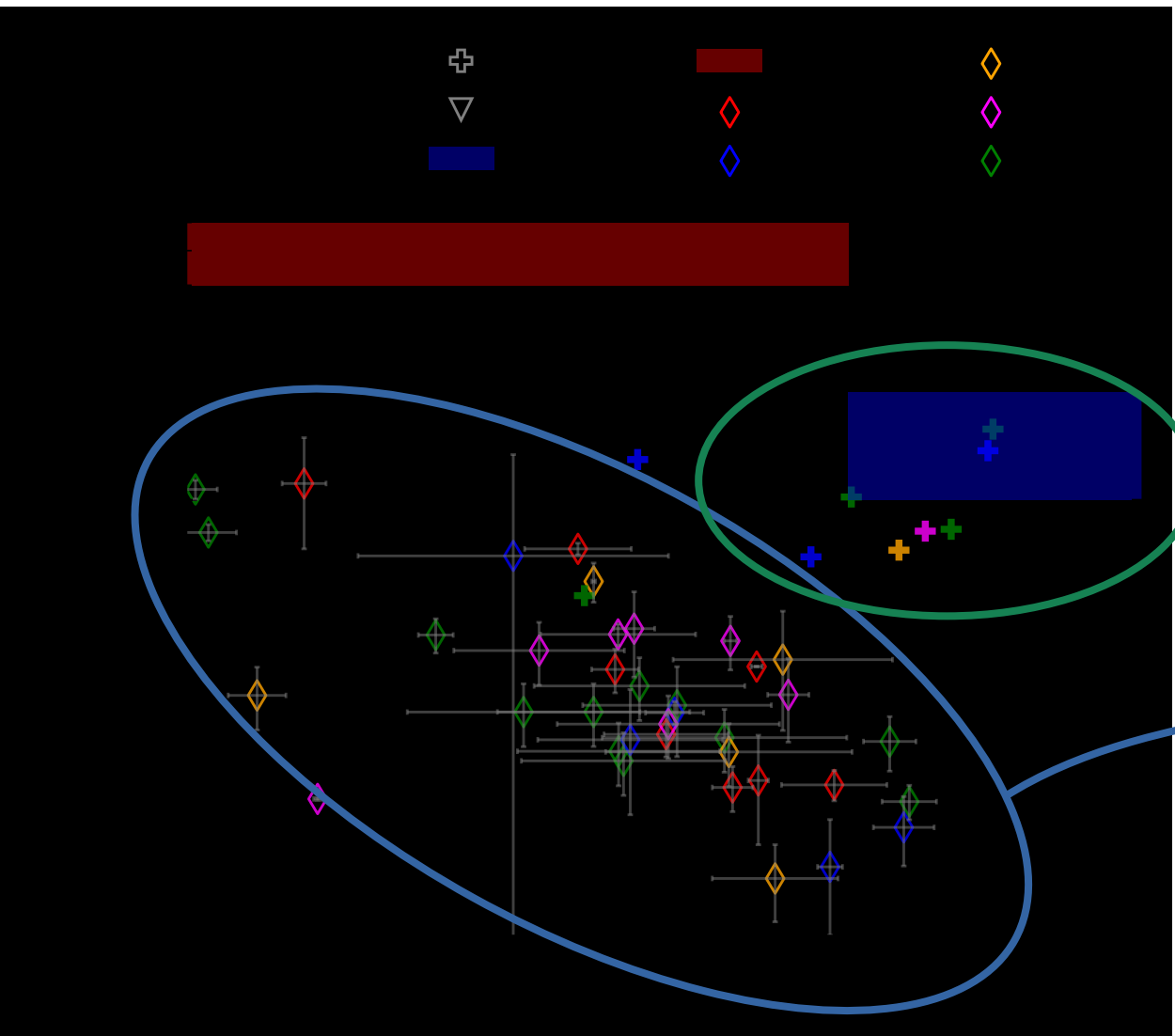
Dynamical ejecta

See also Dessart+ 2009; Fernandez+ 2013; Perego+ 2014, Just+ 2014; ...

From Nedora+, ApJ 906:98 (2021)



Dynamical ejecta

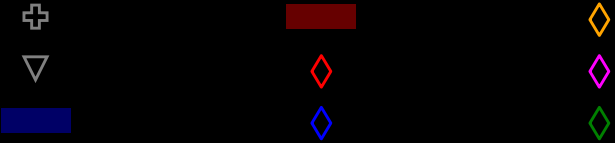


Spiral-wave wind?

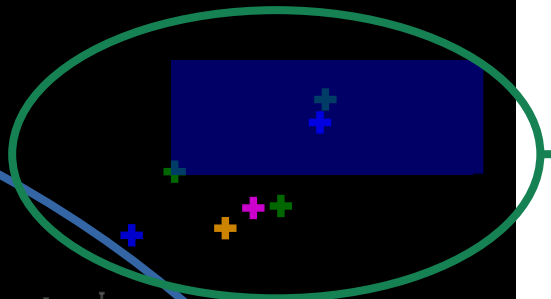
Dynamical ejecta

See also Dessart+ 2009; Fernandez+ 2013; Perego+ 2014, Just+ 2014; ...

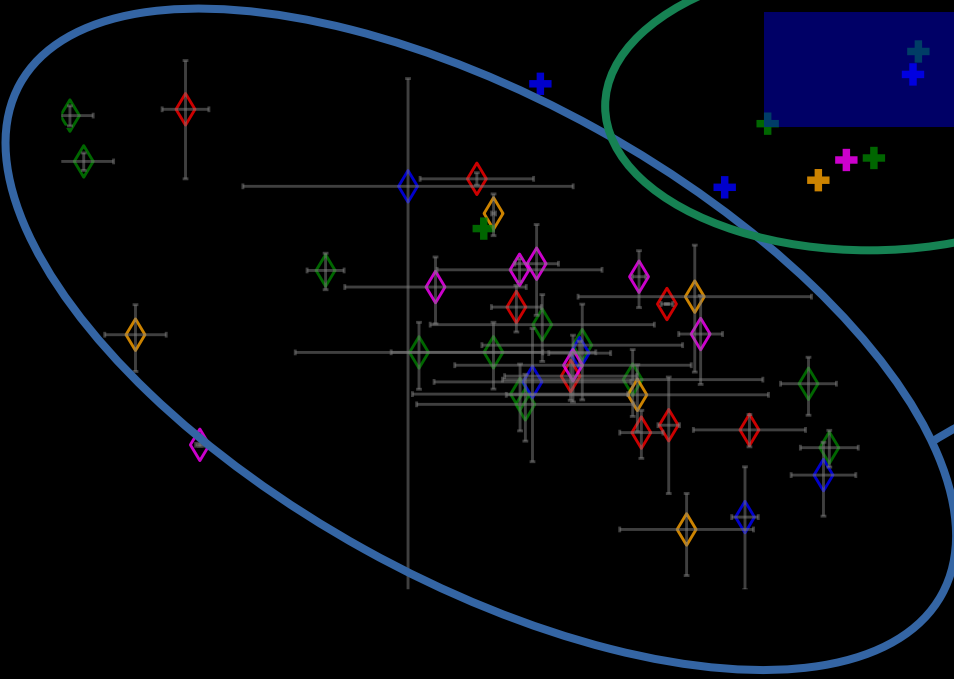
From Nedora+, ApJ 906:98 (2021)



Disk recombination wind?!?

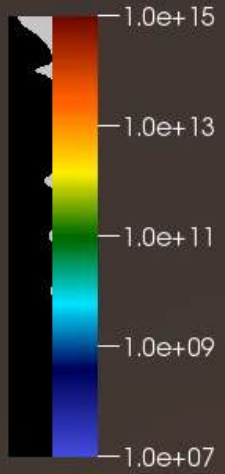


Spiral-wave wind?

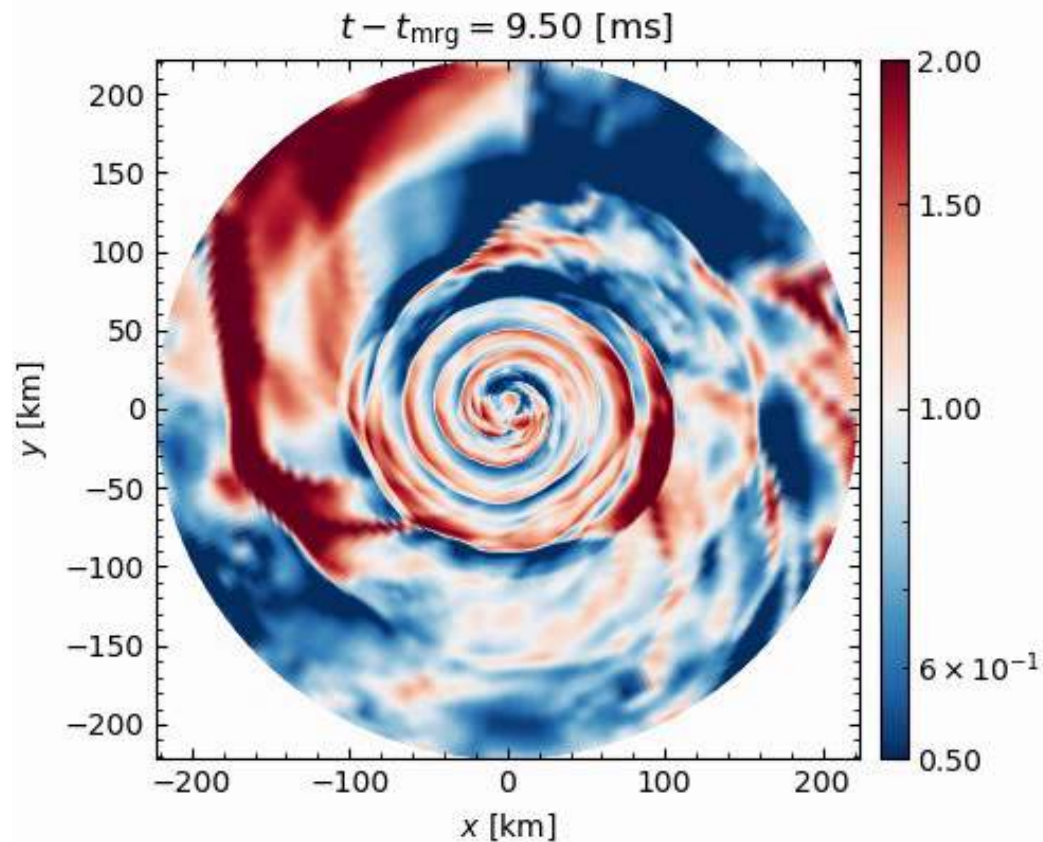
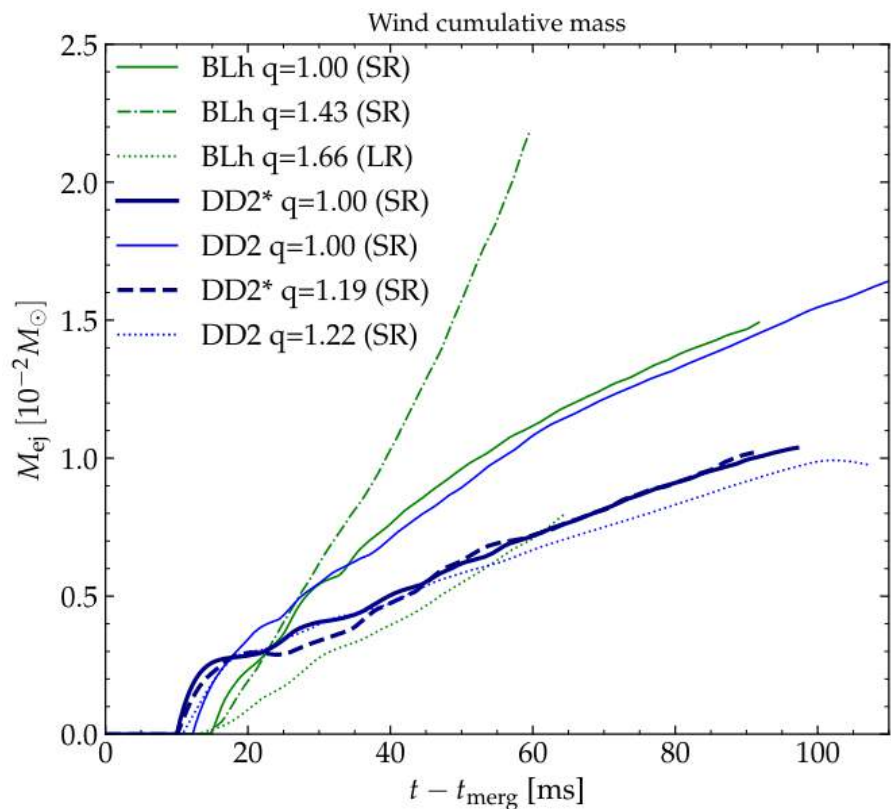


Dynamical ejecta

$t = 1.986 \text{ ms}$



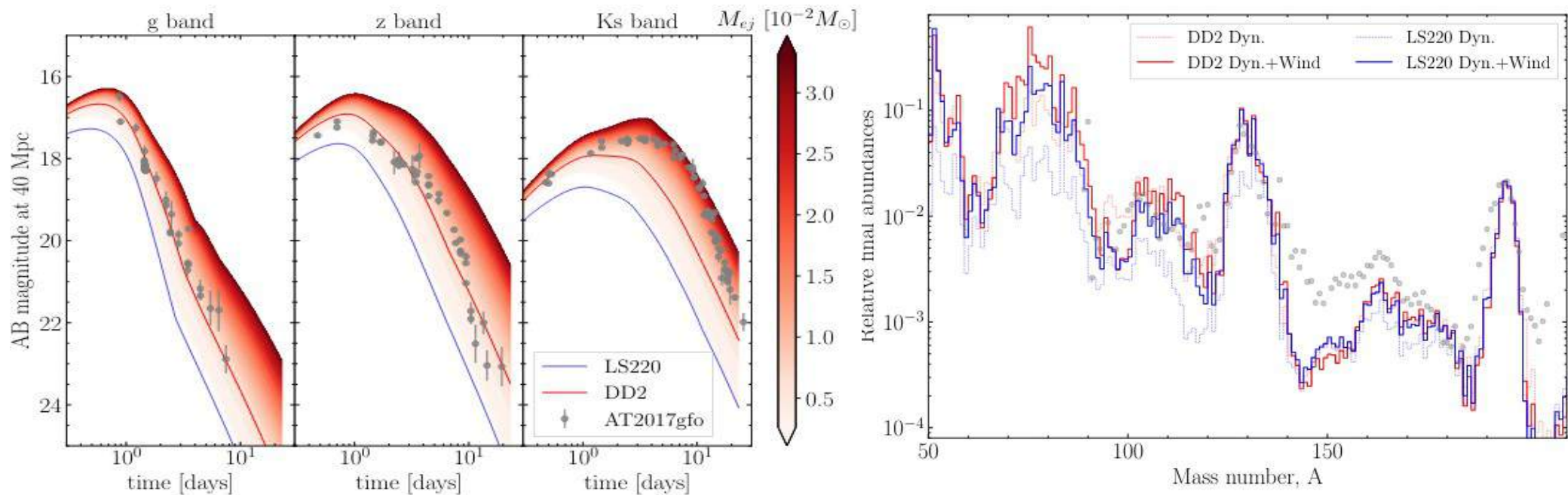
Spiral-wave wind



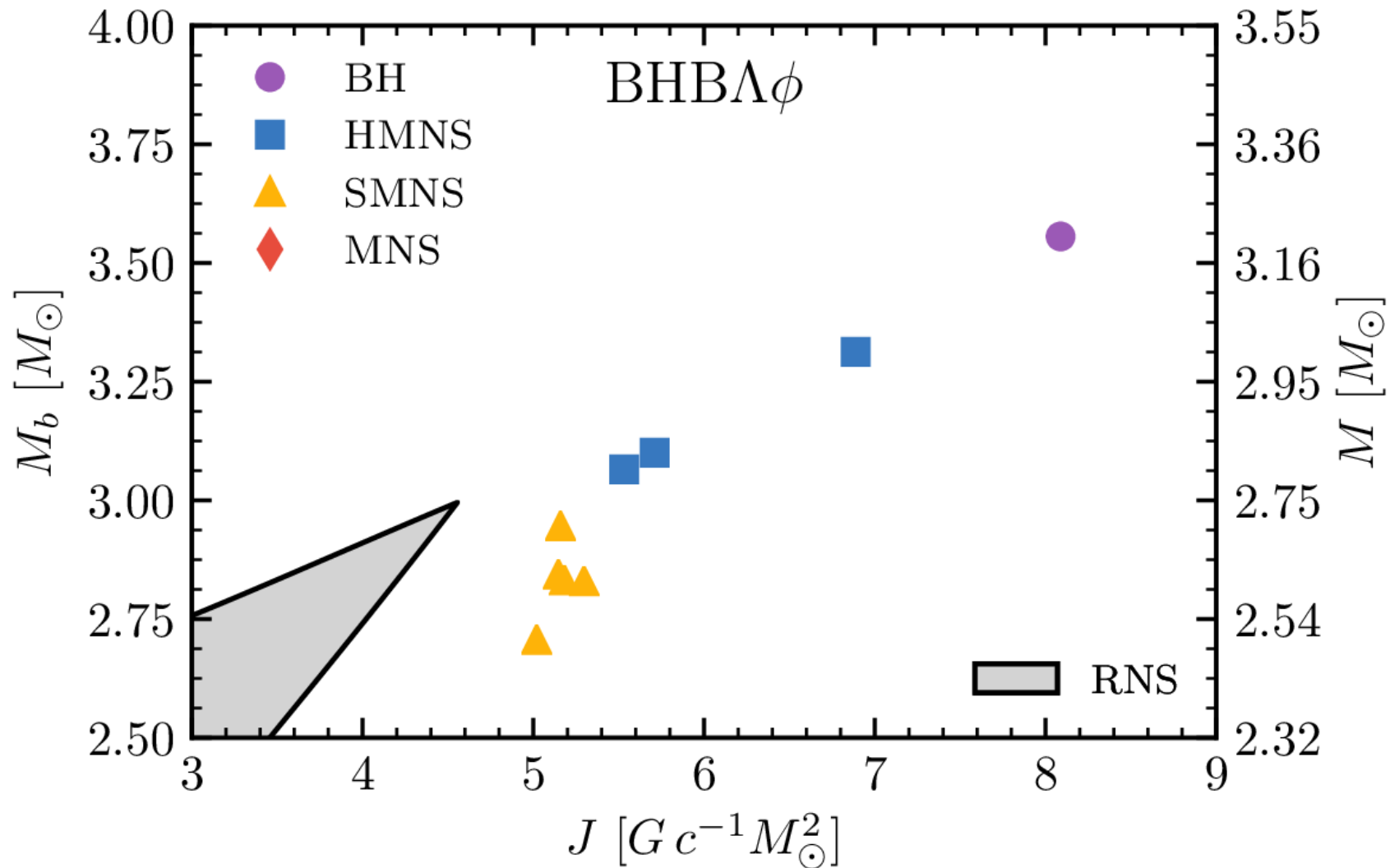
See also Siegel+ 2017; Fujibayashi+ 2018-2023;
Radice+ 2018; Mösta+ 2020; Curtis+ 2023; Combi+ 2023; ...

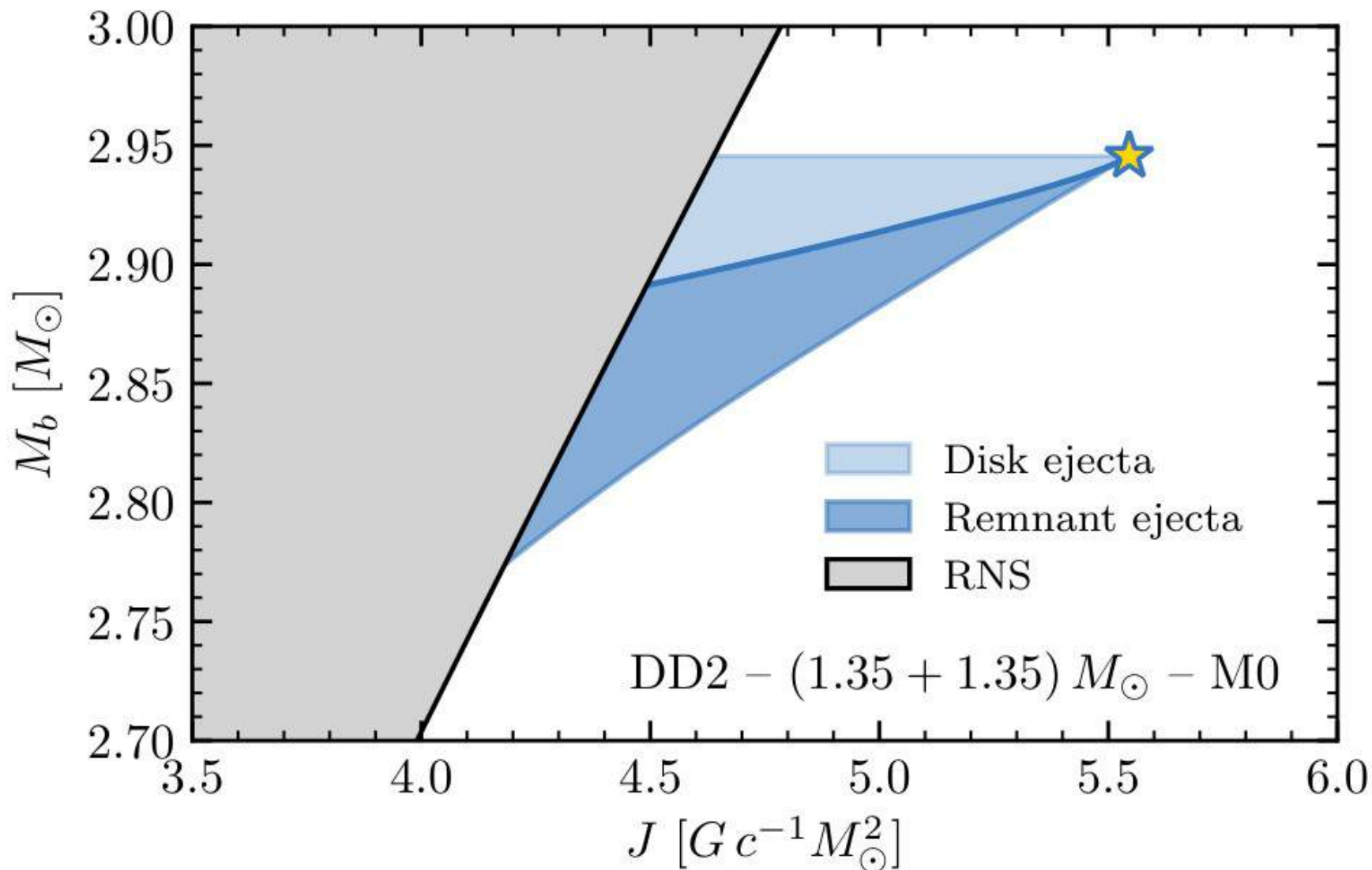
Nedora+ ApJ 906:98 (2021);
Radice+ ApJ 959:46 (2023)

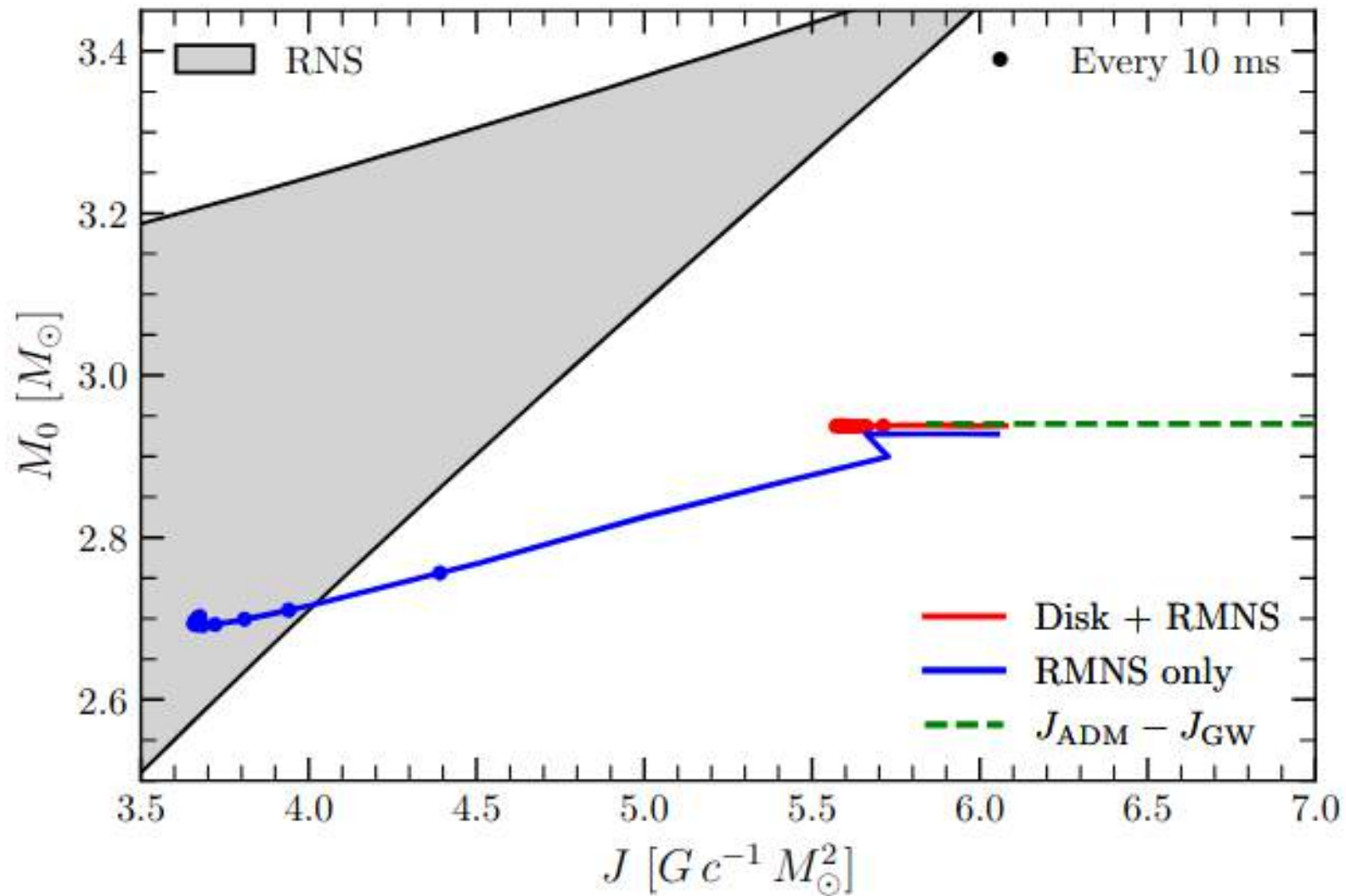
Numerical relativity results



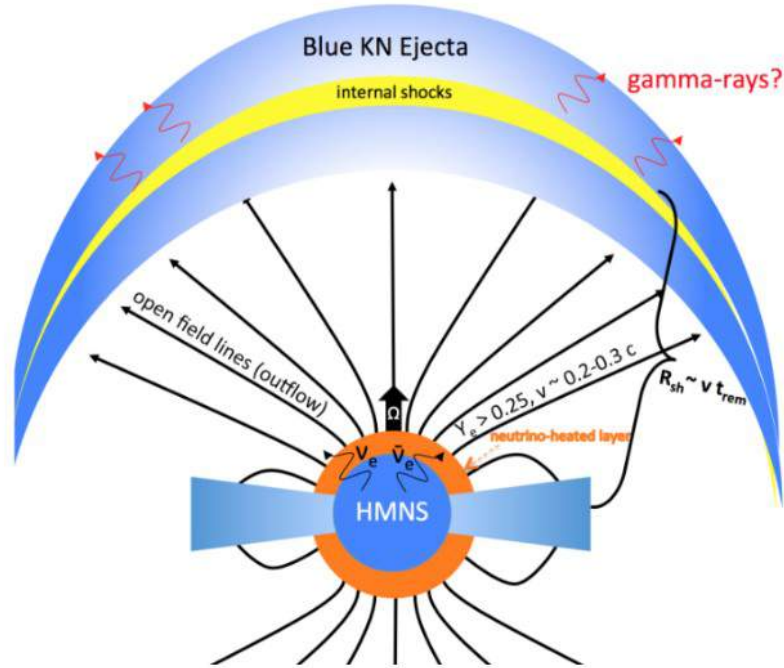
Postmerger winds **naturally** explain blue kilonova emission!



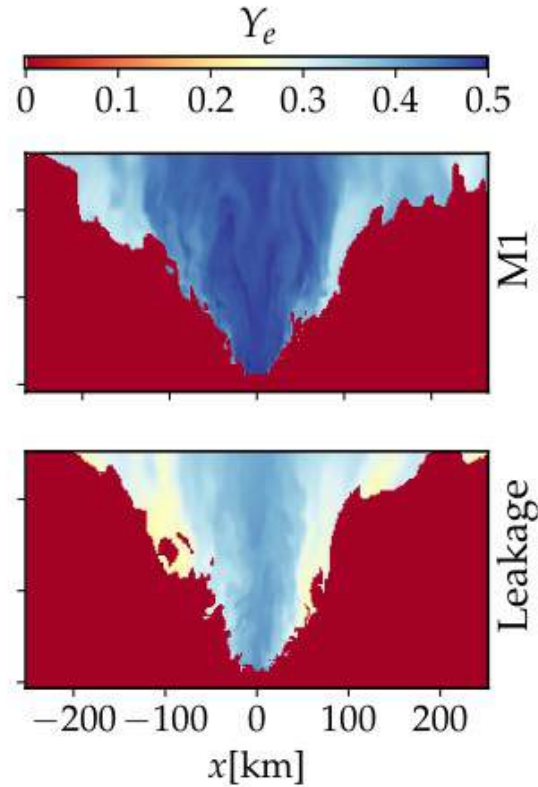




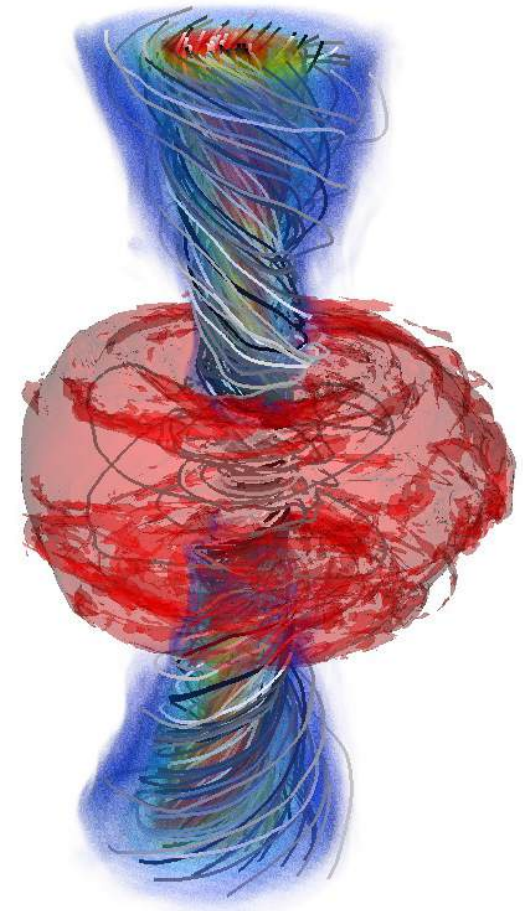
What about magnetic fields?



Metzger+, ApJ 856:101 (2018)

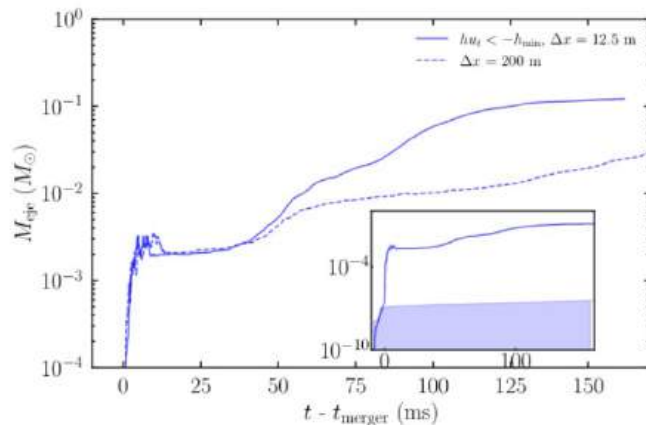
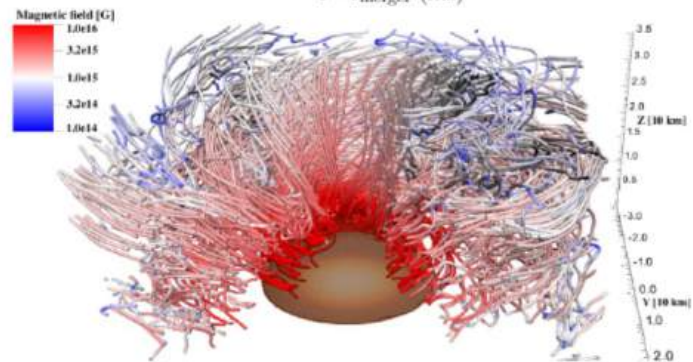
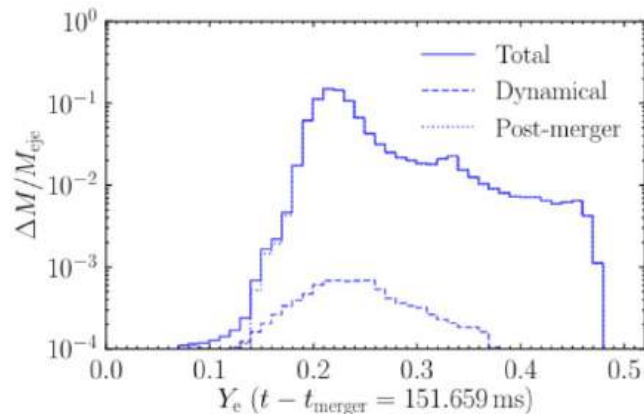
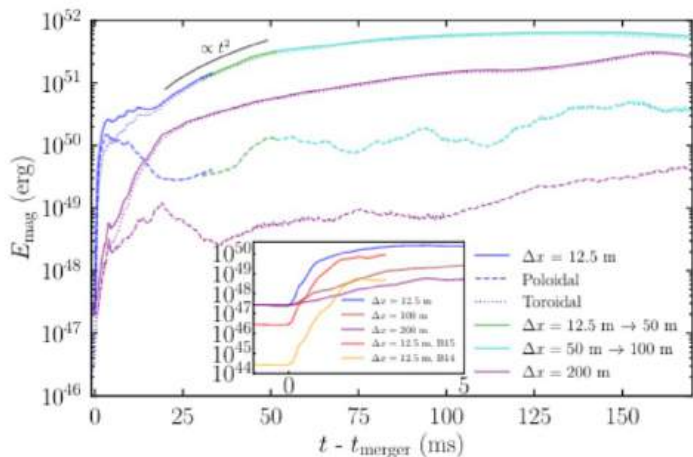


Curtis+, ApJL 961:L26 (2023)

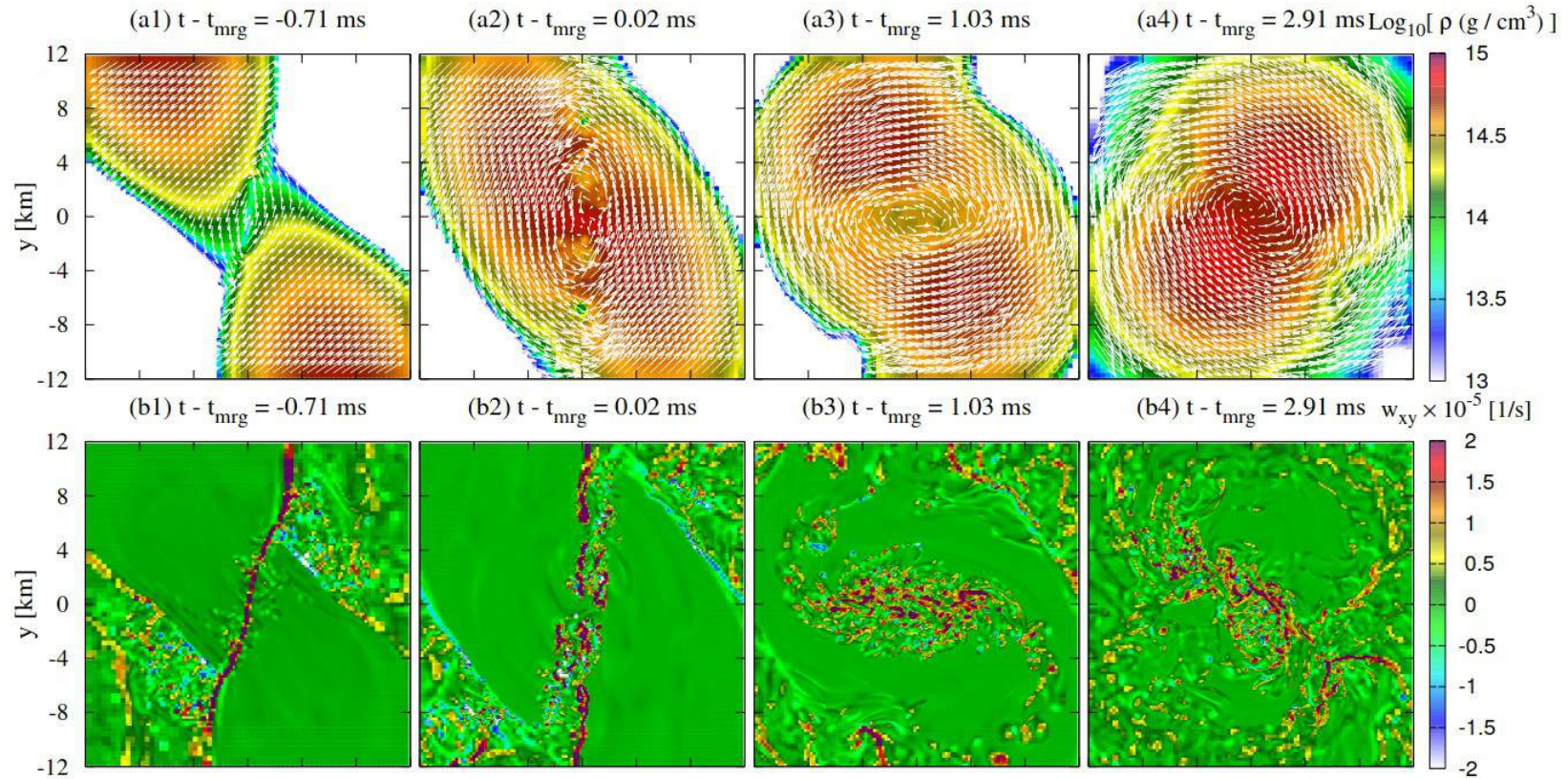


Moesta+, ApJL 901:L37 (2020)

Magnetized winds

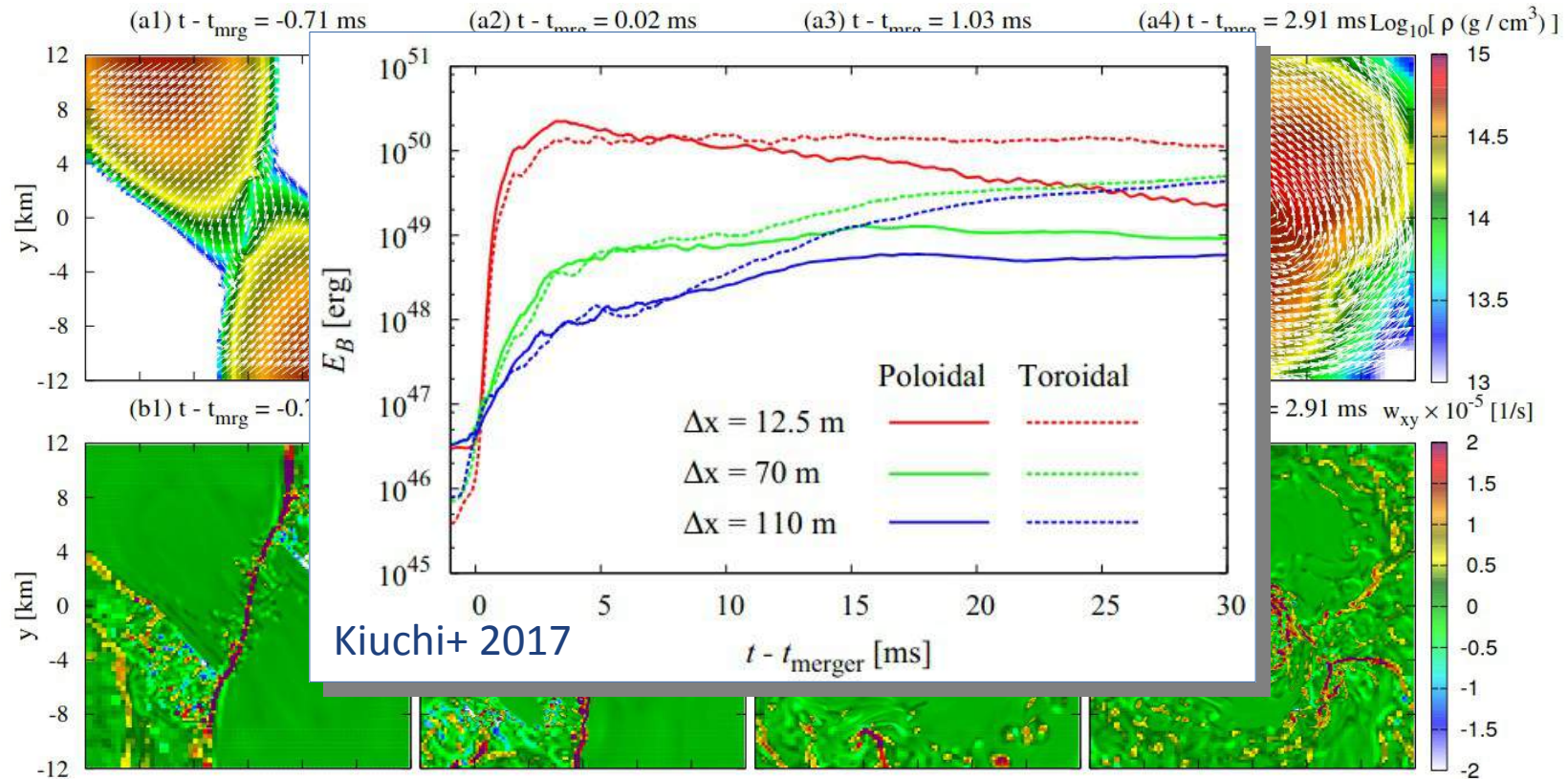


Field amplification by the KHI



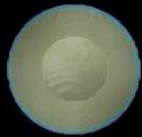
From Kiuchi et al. (2015)

Field amplification by the KHI

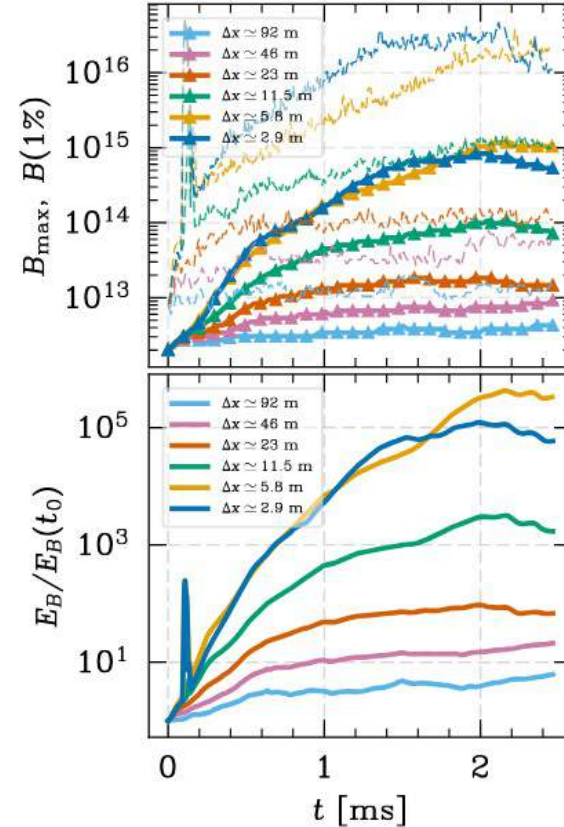
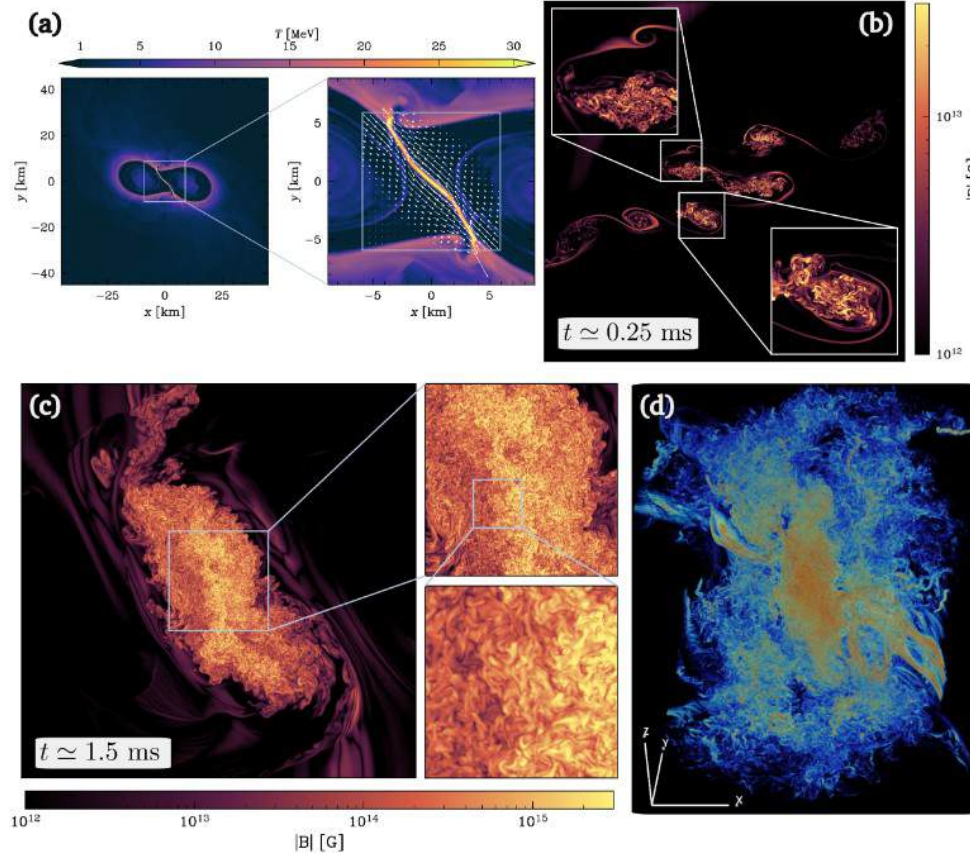


From Kiuchi et al. (2015)

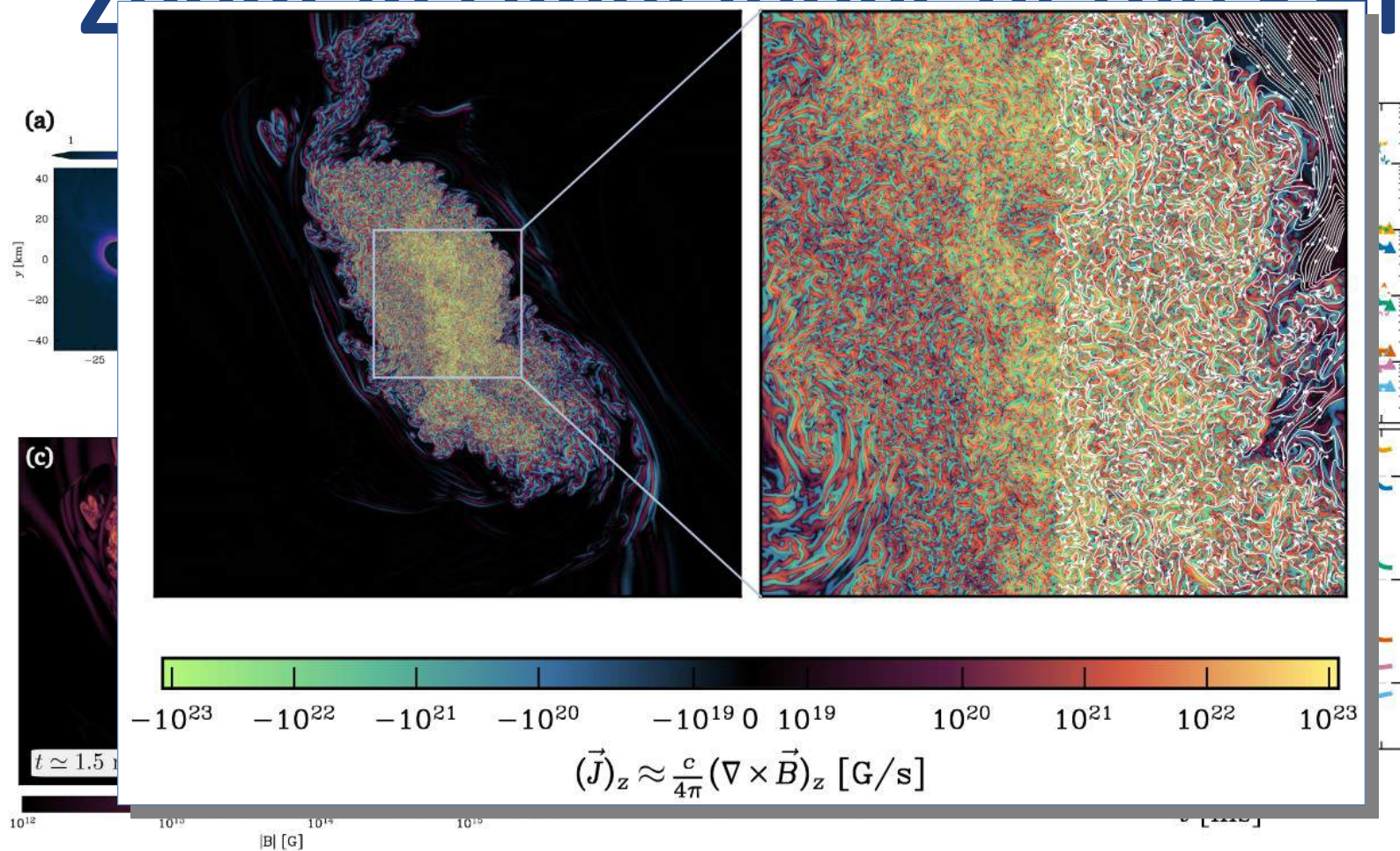
$t = 0.00$ ms



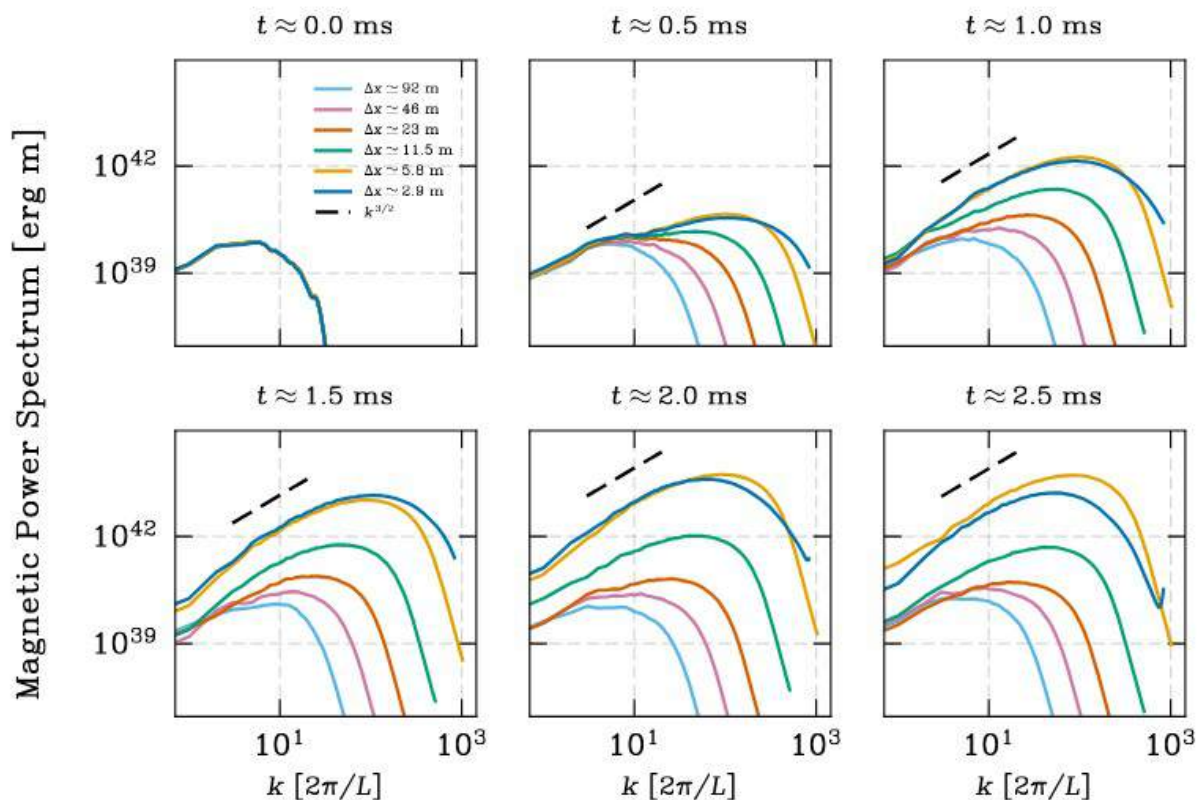
Zoom-in simulations of the KHI



Zoom in simulations of the KHI



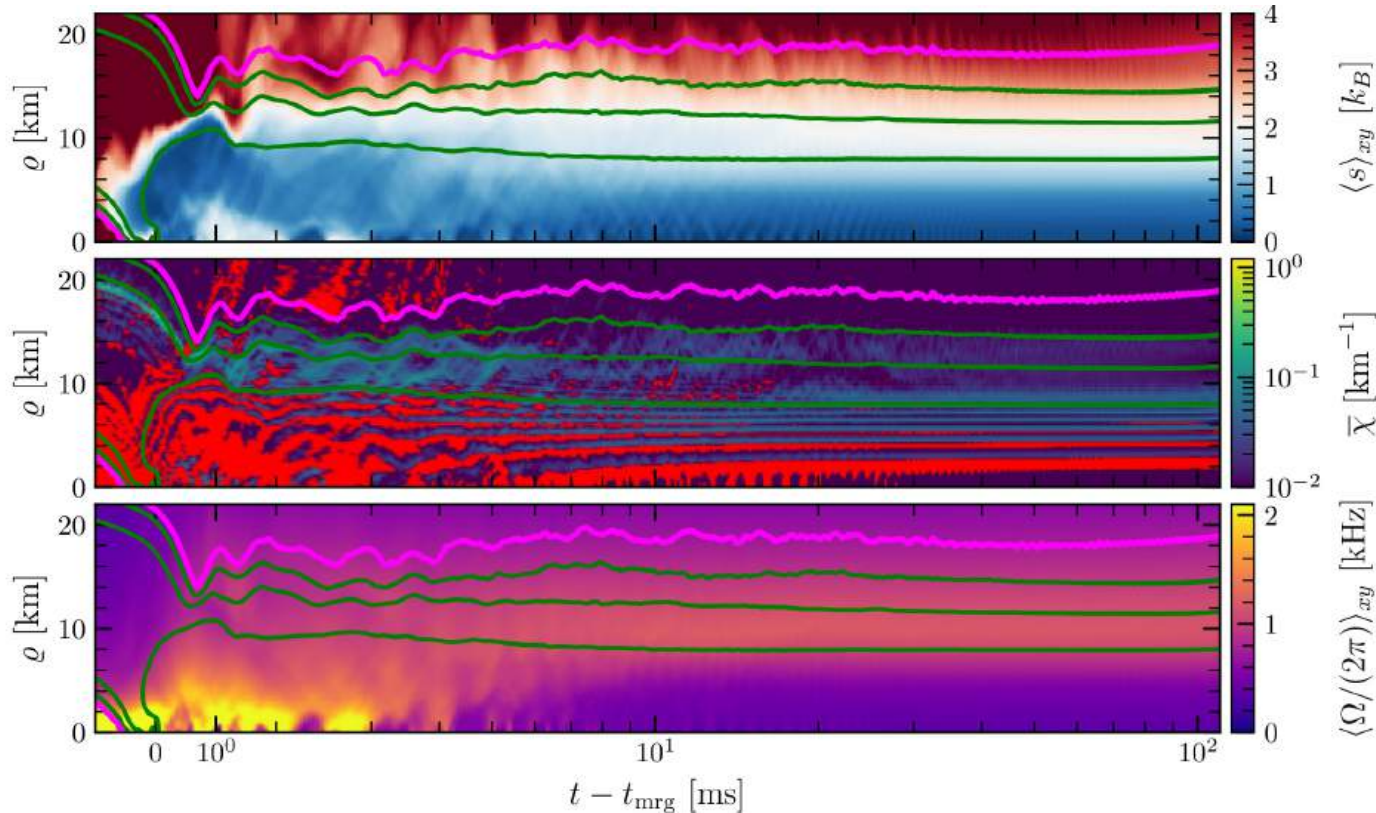
Turbulent dynamo action



See also Price & Rosswog 2006; Obergaulinger+ 2010;
Zrake+2013 Kiuchi+ 2014, 2025; ...

From Gutierréz+, arXiv:2601.20953

Maybe we can bury the field?



From Radice & Bernuzzi, ApJ 959:46 (2023); 995:229 (2025) (erratum)

Flux emergence in NS mergers

$$B_{\max} = 7 \times 10^{17} \text{ G}$$



Time = 0.000 ms



Time = 0.394 ms

$$B_{\max} = 1 \times 10^{17} \text{ G}$$



Time = 0.000 ms



Time = 0.862 ms

Flux emergence in NS mergers

$B_{\max} =$

- Most of the field is buried within the remnant
- Some of the field will leak as the disk is formed. How much?
- Can only rely on the **accretion disk field** to do interesting things

$B_{\max} = 1 \times 10^{17} \text{ G}$

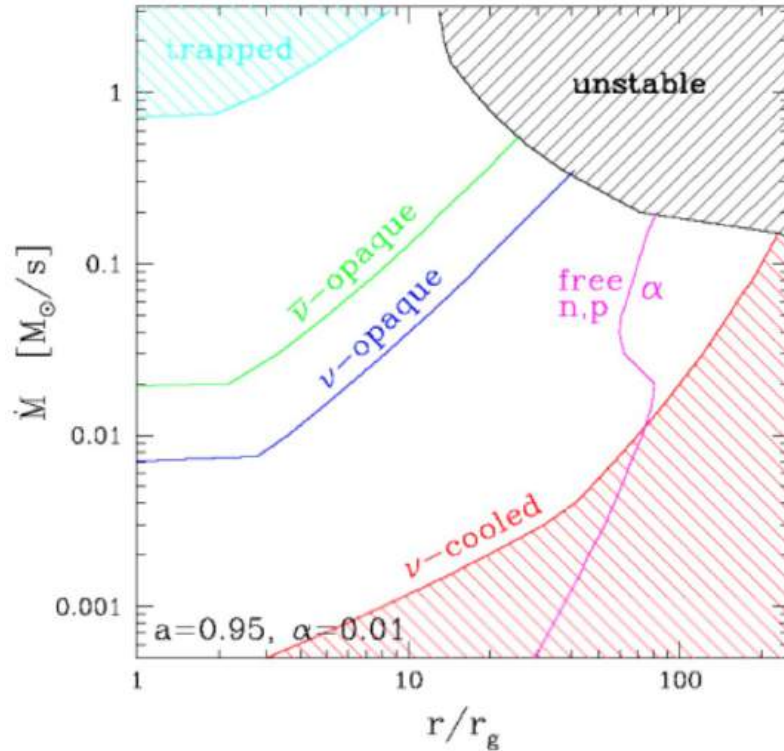


Time = 0.000 ms

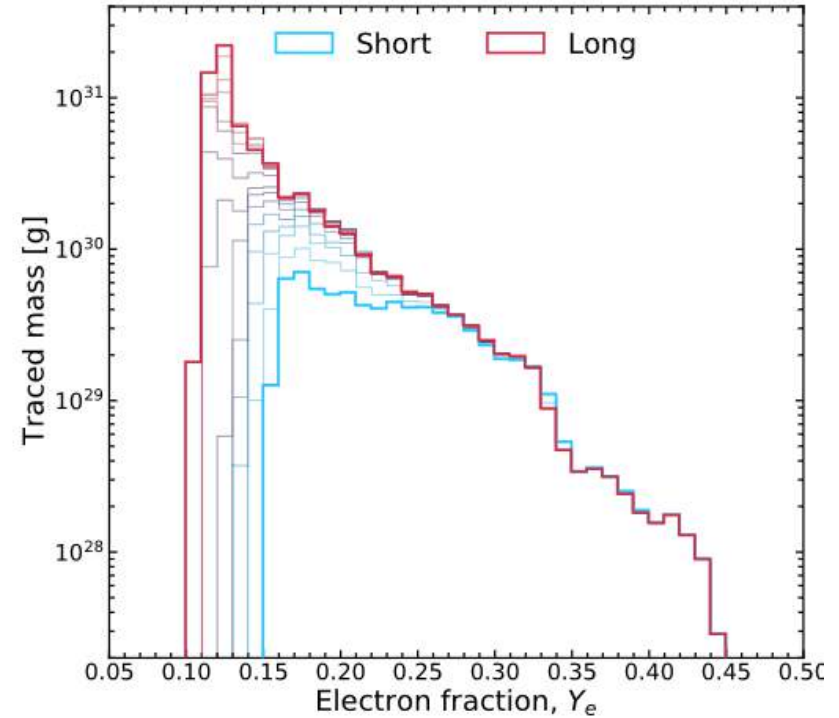


Time = 0.862 ms

Late times: disk evaporation



From Beloborodov AIP Conf.Proc. 1054 (2008)

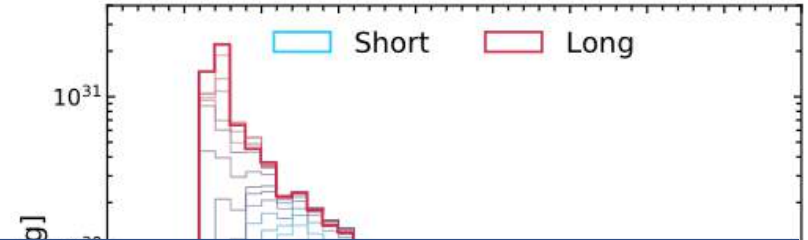
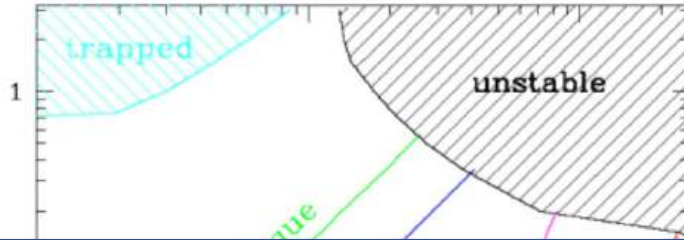


From Sprouse+ ApJ 962:79 (2024)

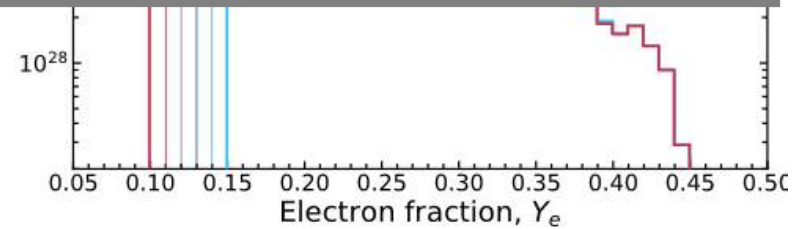
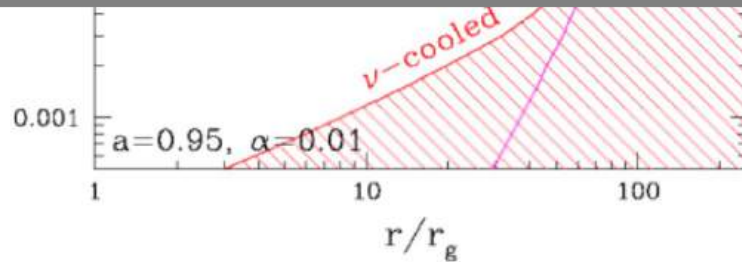
See also Beloborodov (2003); Fernandez+ (2013); Metzger+ (2014); Perego+ (2014); Siegel+ (2017); ...

See Oliver's talk!

Late times: disk evaporation



No **ab-initio** simulation has been able to reproduce the red kilonova in AT2017gfo



From Beloborodov AIP Conf.Proc. 1054 (2008)

From Sprouse+ ApJ 962:79 (2024)

See also Beloborodov (2003); Fernandez+ (2013); Metzger+ (2014); Perego+ (2014); Siegel+ (2017); ...

See Oliver's talk!